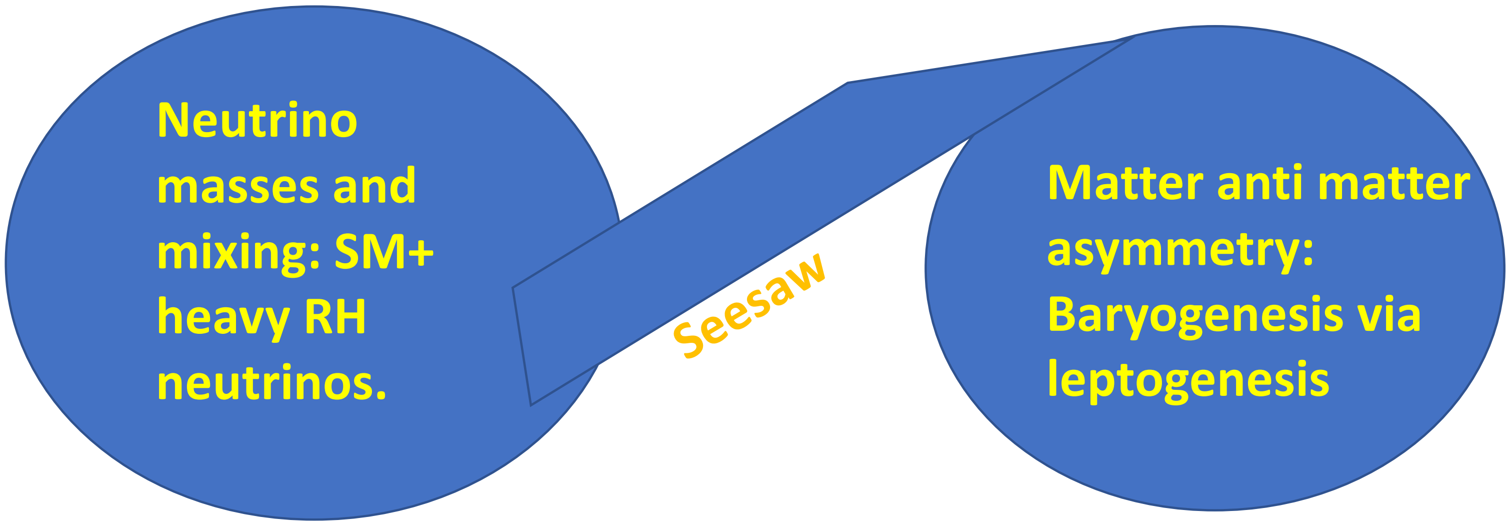
Seesaw neutrino models and their motions in lepton flavor space

Rome Samanta, University of Southampton

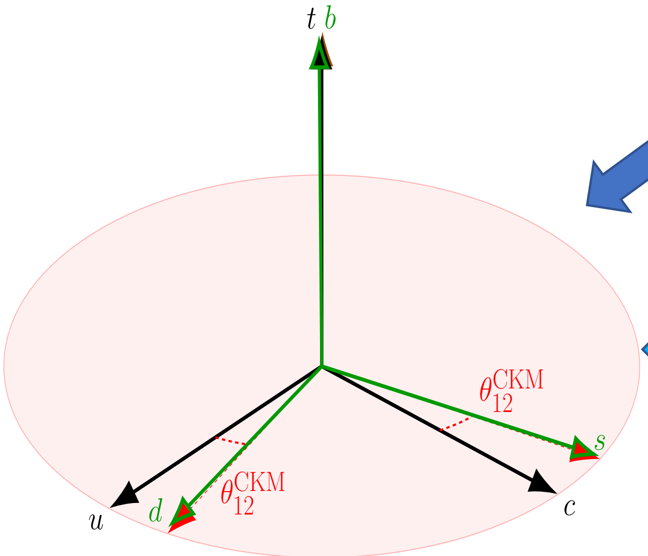


Based on: **`Representing seesaw neutrino models and their motions in lepton flavor space'**, Pasquale Di Bari, Michele Re Fiorentin and **Rome Samanta**, 1812.07720

Neutrino masses and mixing

Figures: P. Di Bari, M. Fiorentin, **RS** Arxiv: 1812.07720

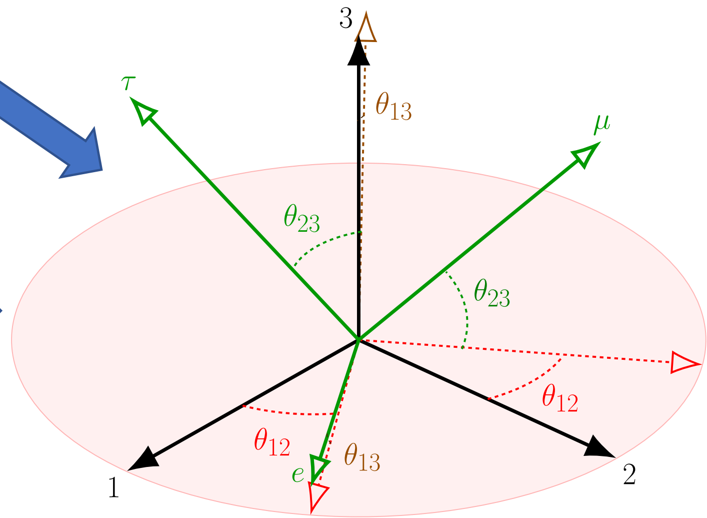
Quarks



GUT ??

Why such mismatch ?

Neutrinos



$\theta_{12}^{CKM} \simeq 13^0$, $\theta_{13}^{CKM} \simeq 0^0$, $\theta_{23}^{CKM} \simeq 0^0$ 🤔 $\theta_{12}^{NEU} \simeq 34^0$, $\theta_{13}^{NEU} \simeq 9^0$, $\theta_{23}^{NEU} \simeq 45^0$

Things to note:

- ✍ **Neutrino oscillation** 👉 Neutrinos have masses.
- ✍ **Cosmology (PLANCK)** 👉 Neutrinos are light, even less than 1 eV.
- ✍ **Standard Model (SM) of particle physics** 👉 cannot explain neutrino masses and mixing.
- ✍ **Need extension of the SM** 👉 Minimal extension requires at least two heavy right handed (RH) neutrinos to explain small neutrino masses trough seesaw mechanism.
- ✍ **No conclusive evidence for antimatter** 👉 AMS experiment is searching for that.
- ✍ **CMB acoustic peak and light elements abundances after BBN** 👉 baryon to photon ratio $\simeq$ **6.2×10^{-10}**
- ✍ Seesaw is a simple and excellent mechanism to explain the baryon asymmetry

Neutrino oscillation data and other cosmological constraints:

Hint for CP violation ($\delta_{CP} = 215^0$) and normal mass ordering ($m_3 > m_2 > m_1$)

$\mathbf{Bf}^{+1\sigma}_{-1\sigma}$ NuFiT, NoV, 2018

$$\theta_{12}^{NEU} \simeq 33.82^{+0.78}_{1-0.762}$$

$$\theta_{13}^{NEU} \simeq 8.61^{+0.13}_{-0.13}$$

$$\theta_{23}^{NEU} \simeq 49.61^{+1.0}_{-1.2}$$

Angles

$$\Delta m_{12}^2 \simeq 7.39^{+0.21}_{-0.20} \times 10^{-5}\ \mathrm{eV}^2$$

$$\Delta m_{23}^2 \simeq 2.52^{+0.033}_{-0.0.032} \times 10^{-3}\ \mathrm{eV}^2$$

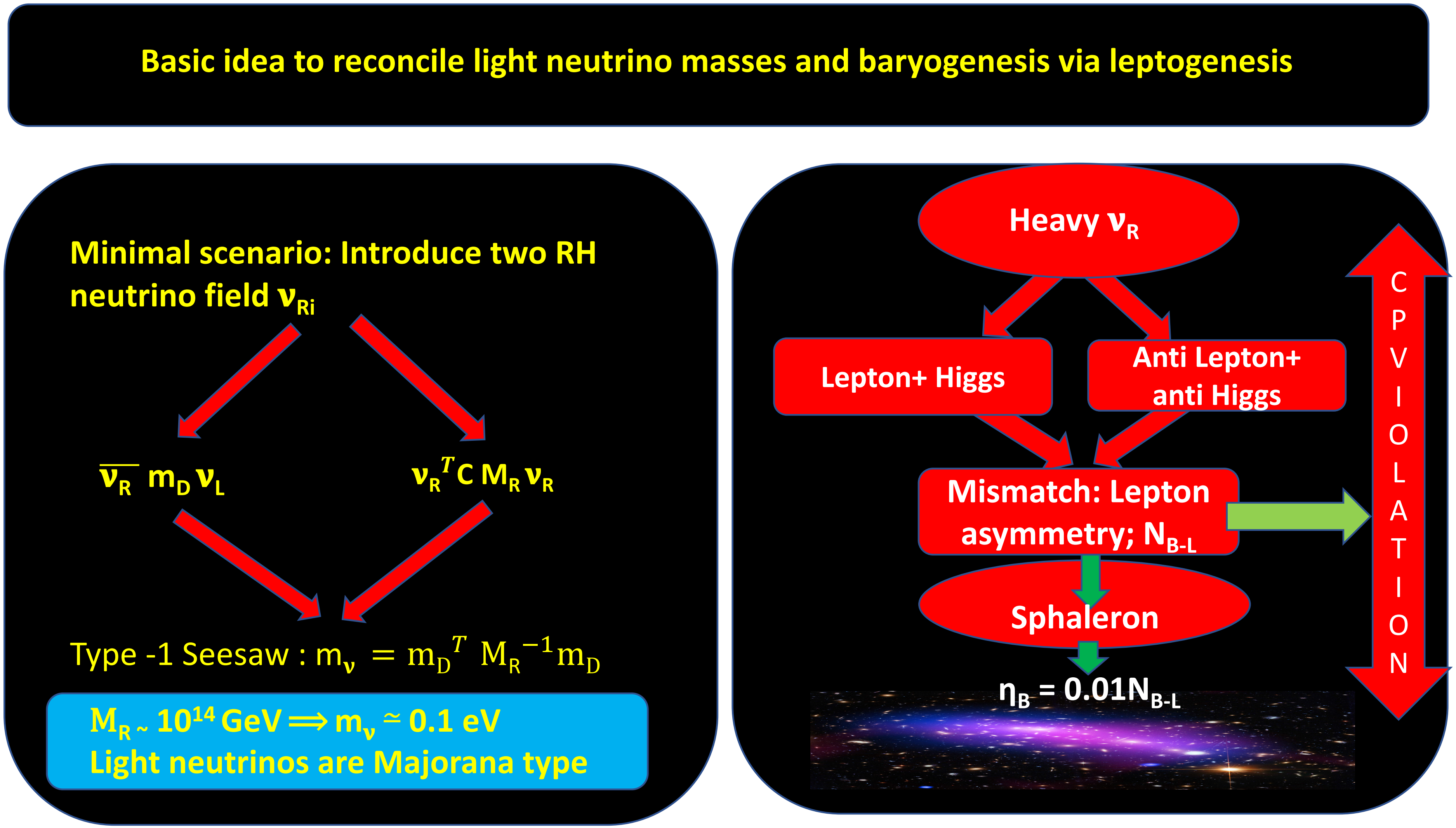
$\sum_i m_i < 0.17$ eV **PLANCK, 2017**

masses

No evidence of antimatter !

- Our Solar system is made of matter
- Even if we think of an antimatter galaxy, we should observe a steady stream of gamma rays arising due to the interaction of the antimatter galaxy and intergalactic matter cloud. We don't see such radiation.
- AMS is looking for the detection of anti matter (anti helium).

Till date observed baryon asymmetry is perfectly consistent with CMB and BBN.

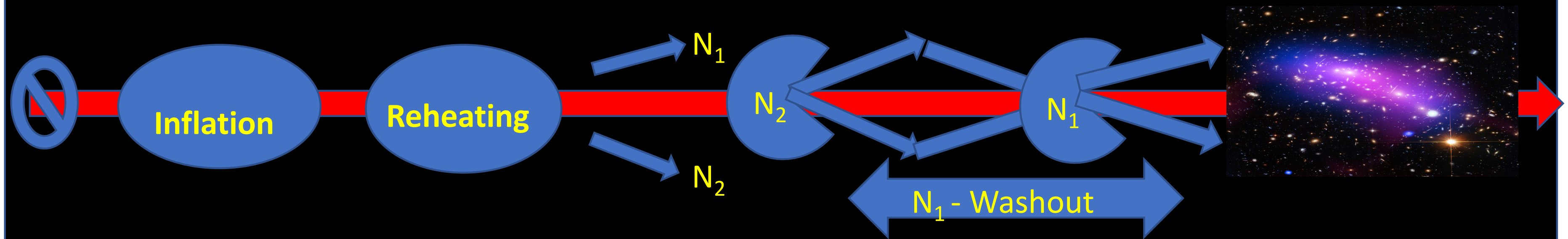


Basic idea to reconcile light neutrino masses and baryogenesis via leptogenesis
Minimal scenario: Introduce two RH neutrino field ν_{Ri}
$\overline{\nu_R}\, m_D\, \nu_L$
$\nu_R^T C\, M_R\, \nu_R$
Type -1 Seesaw : $m_\nu = m_D^T M_R^{-1} m_D$
$M_R \sim 10^{14}$ GeV $\Longrightarrow m_\nu \simeq 0.1$ eV
Light neutrinos are Majorana type
Heavy ν_R
Lepton+ Higgs
Anti Lepton+ anti Higgs
Mismatch: Lepton asymmetry; N_{B-L}
Sphaleron
$\eta_B = 0.01 N_{B-L}$
CP VIOLATION

Types of leptogenesis:

Thermal Leptogenesis:

Compatible with non-SUSY scenario: Gravitino Problem



Non-thermal leptogenesis:

SUSY friendly

RH neutrinos could originate from non-thermal decays of inflaton, compatible with low T_{RH}

The Bridging (B) matrix

Figures: P. Di Bari, M. Fiorentin, RS Arxiv: 1812.07720

$$-\mathcal{L}^{\nu+\ell}_{Y+M} = \overline{L_\alpha}\, h^{\ell}_{\alpha\alpha}\, \ell_{R\alpha}\, \Phi + \overline{L_\alpha}\, h^{\nu}_{\alpha J}\, N_{RJ}\, \widetilde{\Phi} + \frac{1}{2}\, \overline{N^c_{RJ}}\, M_J\, N_{RJ} + \text{h.c.}$$

$$|L_J\rangle = \frac{m_{D\alpha J}}{\sqrt{(m_D^\dagger\, m_D)_{JJ}}}\, |L_\alpha\rangle$$

$$|L_J\rangle = \frac{m_{D\alpha J}\, U^\star_{\alpha i}}{\sqrt{(m_D^\dagger\, m_D)_{JJ}}}\, |L_i\rangle = \frac{(U^\dagger\, m_D)_{iJ}}{\sqrt{(m_D^\dagger\, m_D)_{JJ}}}\, |L_i\rangle$$

$$B_{iJ} \equiv \frac{(U^\dagger\, m_D)_{iJ}}{\sqrt{(m_D^\dagger\, m_D)_{JJ}}}$$

$$p^0_{IJ} \equiv |\langle L_J | L_I \rangle|^2 = \left| \sum_k B^\star_{kJ}\, B_{kI} \right|^2$$

$$B_{iJ} = \sqrt{\frac{m_i}{\widetilde{m}_J}}\, \Omega_{iJ}$$

$$m_{D\alpha J} = U_{\alpha i}\, \sqrt{m_i}\, \Omega_{iJ}\, \sqrt{M_J}\, .$$

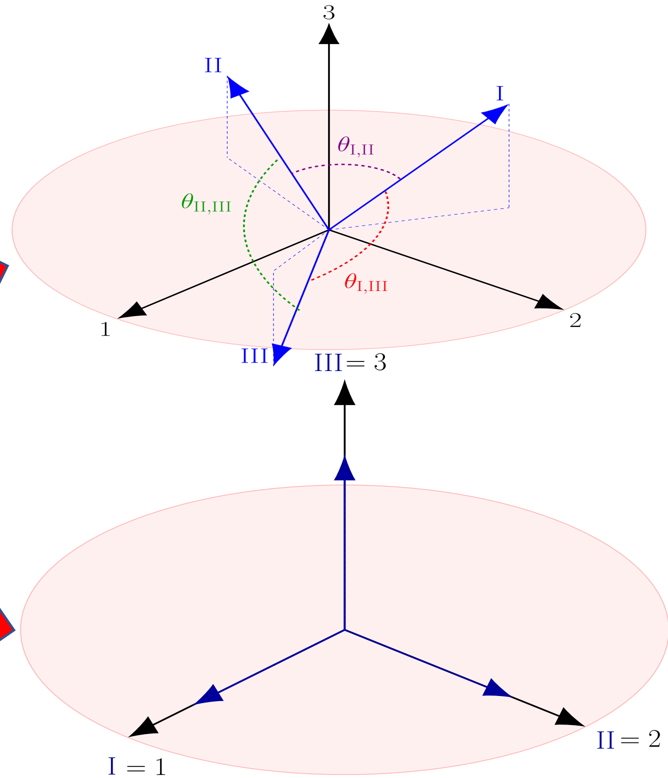
$$\widetilde{m}_J \equiv \frac{(m_D^\dagger\, m_D)_{JJ}}{M_J} = \sum_k m_k\, |\Omega_{kJ}|^2$$

$$m_D^\dagger\, m_D = \lambda_D^2\, P,$$

$$\Omega = P$$

$$m_i = \lambda_D^2 / M_J$$

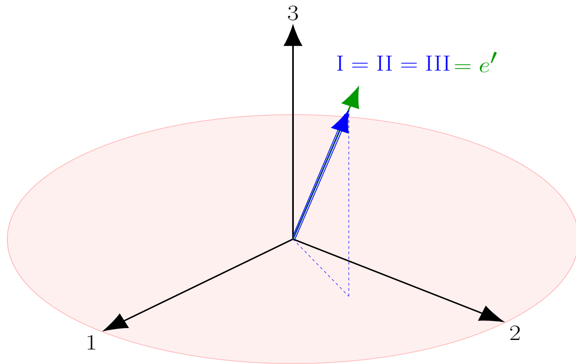
JHEP 0906 (2009) 072
SF King, Mu Chun Chen



Form Dominance

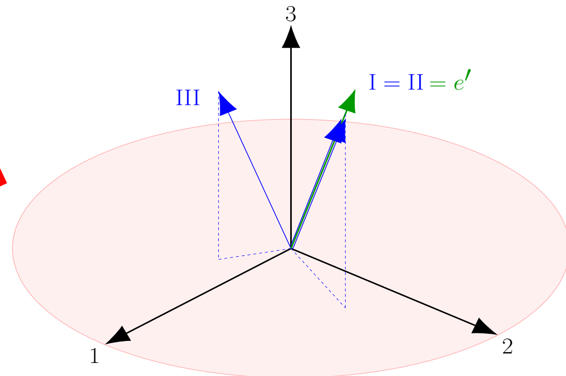
Some examples

$$m'_D = V'_L\, m_D = \begin{pmatrix} m_{De'I} & m_{De'II} & m_{De'III} \\ 0 & 0 & 0 \\ 0 & 0 & 0 \end{pmatrix}$$



$$m'_D = V'_L\, m_D = \begin{pmatrix} m_{De'I} & m_{De'II} & m_{De'III} \\ 0 & 0 & m_{D\mu'III} \\ 0 & 0 & m_{D\tau'III} \end{pmatrix}$$

Creates zeros in U_{PMNS}



Phys. Lett. B 644 (2007) 59 , Mohapatra et al.
JCAP 1703 (2017) no.03, 025 , RS et al.
JHEP 1712 (2017) 030, RS et al.

Fine tuning in the seesaw and a new parametrization of the orthogonal matrix

$$m_i = m_{D\ell}^2 / M_J$$

$$m_i = \overline{m}_i \sum_J r_{iJ}\, e^{i\,\varphi_{iJ}} \,,$$

$$\overline{m}_i \equiv m_i \sum_J |\Omega_{iJ}^2|.$$

$$r_{iJ} \equiv |\Omega_{iJ}^2| / \sum_J |\Omega_{iJ}^2| \propto 1/M_J$$

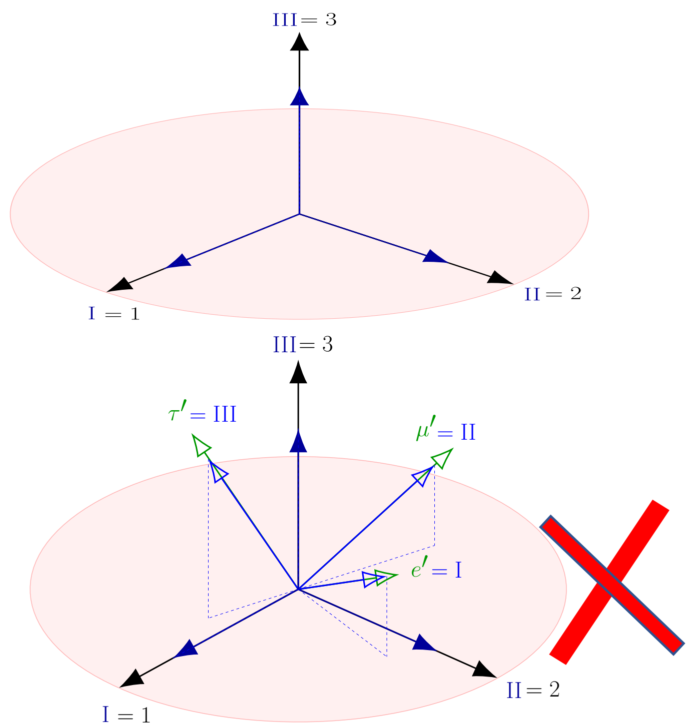
Fine tuning parameter:

$$\gamma_i \equiv \sum_J |\Omega_{iJ}^2| \geq 1$$

$$\Omega = \zeta \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos z_{23} & \sin z_{23} \\ 0 & -\sin z_{23} & \cos z_{23} \end{pmatrix} \begin{pmatrix} \cos z_{13} & 0 & \sin z_{13} \\ 0 & 1 & 0 \\ -\sin z_{13} & 0 & \cos z_{13} \end{pmatrix} \begin{pmatrix} \cos z_{12} & \sin z_{12} & 0 \\ -\sin z_{12} & \cos z_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

$$\Omega(z_{12}, z_{13}, z_{23}) = R(\alpha_{12}, \alpha_{13}, \alpha_{23}) \cdot \Omega_{\rm boost}(\vec{\beta}) \,,$$

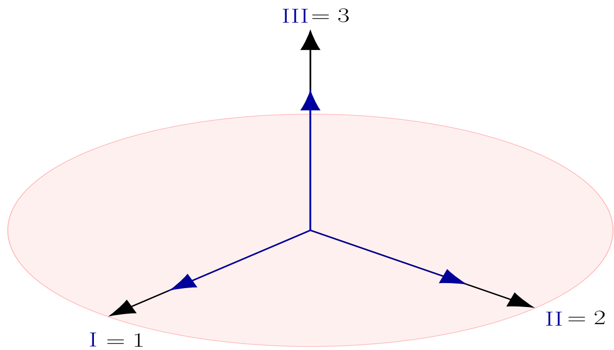
SO(3,C) isomorphic to the proper Lorentz group



A new parametrization for the orthogonal matrix: Lorentz boost in the flavour space

$$\Omega(z_{12}, z_{13}, z_{23}) = R(\alpha_{12}, \alpha_{13}, \alpha_{23}) \cdot \Omega_{\mathrm{boost}}(\vec{\beta}) \,,$$

R is the usual SO(3) rotation matrix



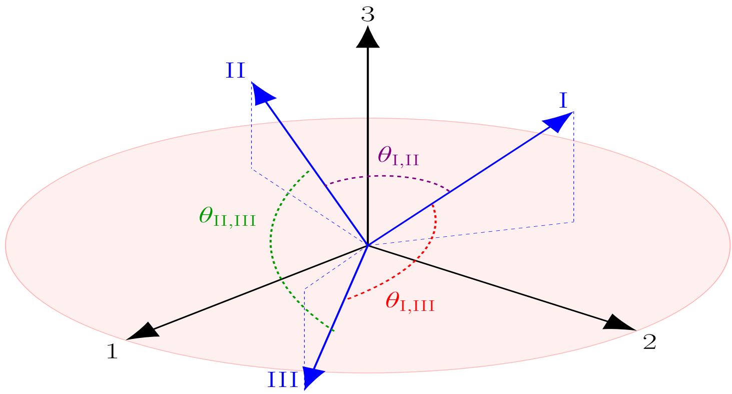
$$\Omega^{\mathrm{Boost}}(\xi, \hat{n}) = \begin{pmatrix} \cosh\xi + n_1^2(1-\cosh\xi) & n_1 n_2(1-\cosh\xi) - i n_3 \sinh\xi & n_1 n_3(1-\cosh\xi) + i n_2 \sinh\xi \\ n_1 n_2(1-\cosh\xi) + i n_3 \sinh\xi & \cosh\xi + n_2^2(1-\cosh\xi) & n_2 n_3(1-\cosh\xi) - i n_1 \sinh\xi \\ n_1 n_3(1-\cosh\xi) - i n_2 \sinh\xi & n_2 n_3(1-\cosh\xi) + i n_1 \sinh\xi & \cosh\xi + n_3^2(1-\cosh\xi) \end{pmatrix}$$

$$\Omega_{\mathrm{boost}}(0,0,\beta) = \begin{pmatrix} \cosh\psi & -i\sinh\psi & 0 \\ i\sinh\psi & \cosh\psi & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

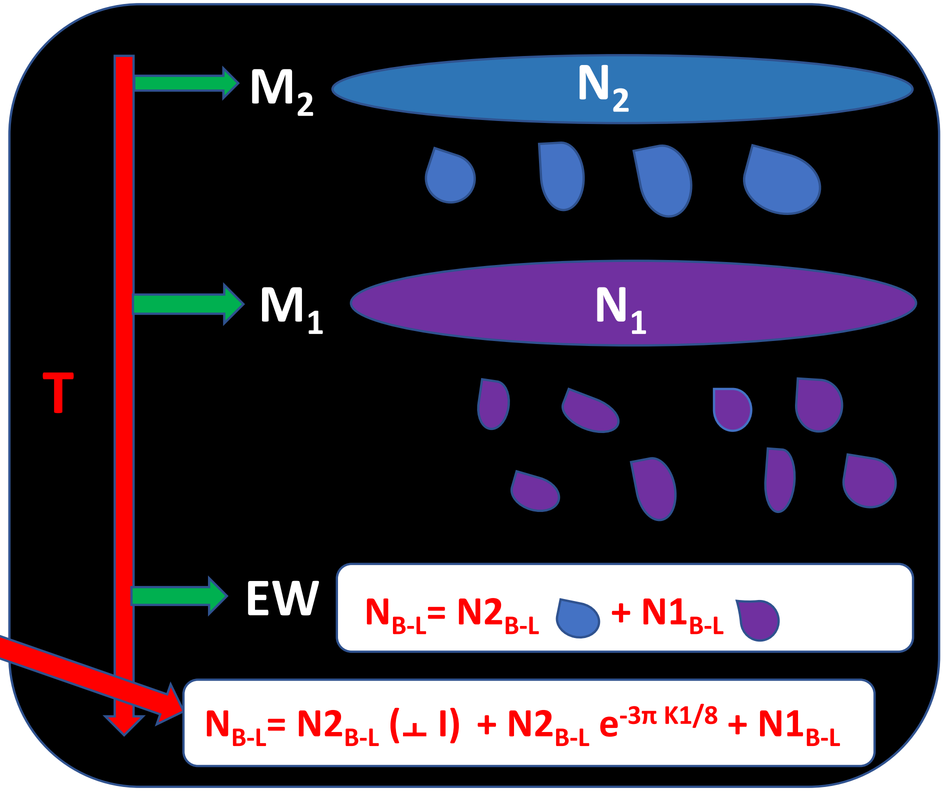
$$\gamma_i \equiv \sum_J |\Omega_{iJ}^2| \geq 1$$

$$\gamma_1 = \gamma_2 = \gamma^2 \left(1+\beta^2\right)$$

One flavour leptogenesis : Computation of the lepton asymmetry

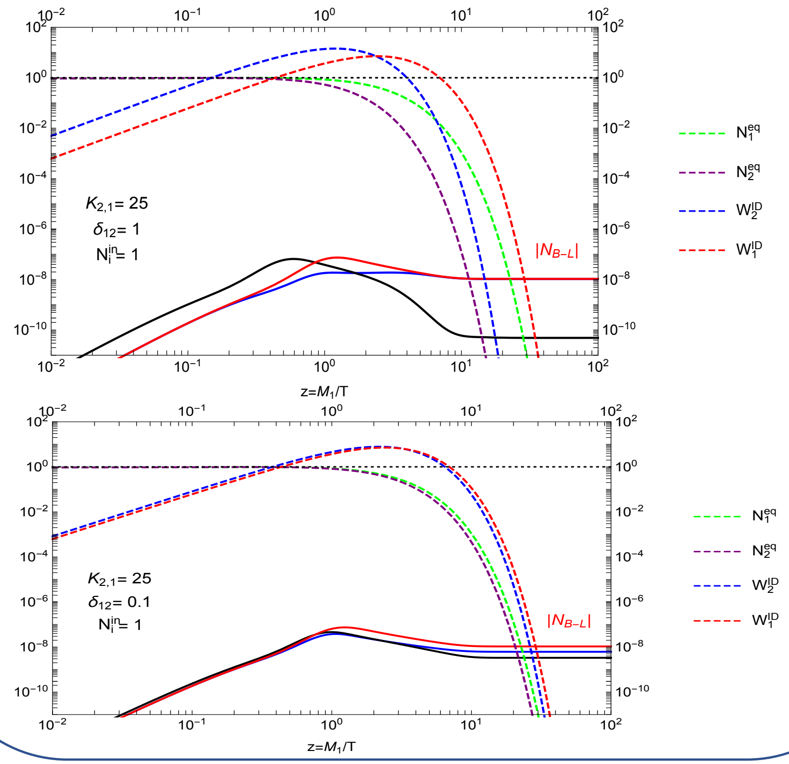


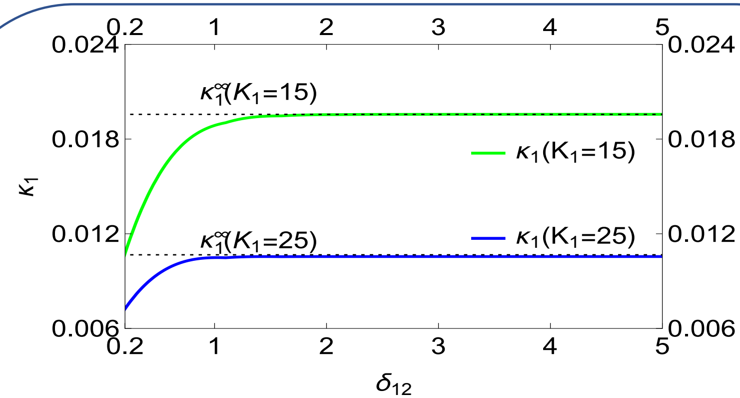
N_1 can only washout the asymmetry generated by N_2 in the direction of $\vec{I}$. Component orthogonal to $\vec{I}$ will always survive. Hence there will always be a survival asymmetry generated by N_2 except in a special case where $\Theta_{I,II} = 0$.

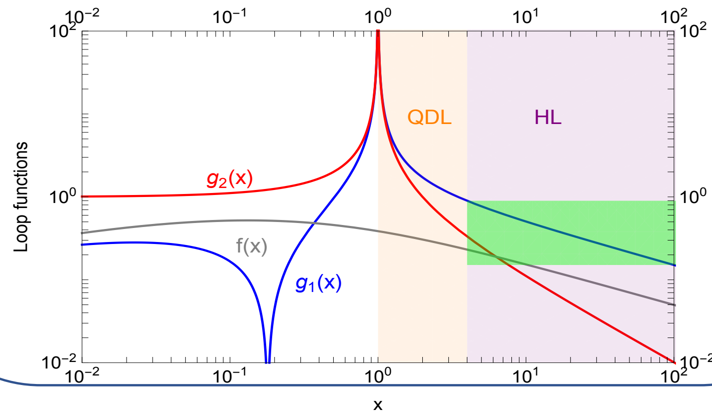


One flavor leptogenesis : Computation of the lepton asymmetry

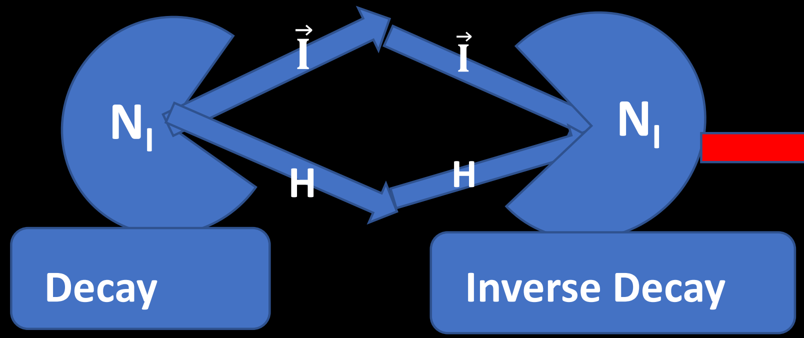
Caution: We are only discussing the hierarchical scenario

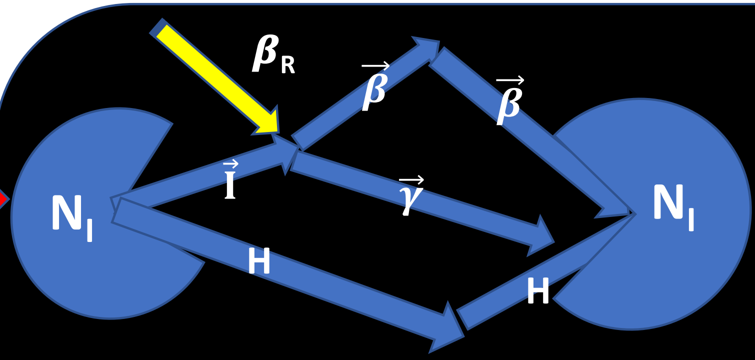






Importance of flavor effects:





If 10^9 GeV < M_l < 10^{12} GeV: $\vec{\beta} = \vec{\tau}$, $\vec{\gamma} = \vec{e} + \vec{\mu}$

$M_l < 10^9$ GeV : All the three flavors act individually

$\vec{l} = P_{l\alpha}\vec{\alpha}$, $\alpha = e, \mu, \tau$, $P_{l\alpha} = K_{i\alpha} / \Sigma_\alpha K_{i\alpha}$

$K_{i\alpha} = P_{l\alpha} K_i$ ← Magnitude of the Decay parameter has been reduced

One flavor leptogenesis : Computation of the lepton asymmetry

(sorry for showing so many equations!)

Boltzmann Equations:

$$\frac{dN_i}{dz} = -D_i(N_{N_i} - N_{N_i}^{\rm eq}), \text{ with } i = 1,2$$

$$\frac{dN_{B-L}}{dz} = -\sum_{i=1}^{2} \varepsilon_i D_i(N_{N_i} - N_{N_i}^{\rm eq}) - \sum_{i=1}^{2} W_i N_{B-L},$$

Inverse Decay:

$$W_i^{\rm ID} = \frac{1}{4} K_i \sqrt{x_{1i}} \mathcal{K}_1(z_i) z_i^3.$$

$Z_i = M_i/T$, $x_{1i} = (M_i/M_1)^2$

$$\kappa_i(z) = -\int_{z_{\rm in}}^{\infty} \frac{dN_{N_i}}{dz'} e^{-\sum_i \int_{z'}^{z} W_i^{\rm ID}(z'')dz''} dz'.$$

$$\kappa_1^\infty = \frac{2}{K_1 z_B(K_1)} \left(1 - e^{-\frac{K_1 z_B(K_1)}{2}}\right),$$

$$\kappa_2^\infty = \frac{2}{K_2 z_B(K_2)} \left(1 - e^{-\frac{K_2 z_B(K_2)}{2}}\right) e^{-\int_0^\infty W_1^{\rm ID}(z)dz}$$

$$\equiv \frac{2}{K_2 z_B(K_2)} \left(1 - e^{-\frac{K_2 z_B(K_2)}{2}}\right) e^{-3\pi K_1/8},$$

where

$$z_B(K_i) = 2 + 4K_i^{0.13} e^{-\frac{2.5}{K_i}}$$

and one uses

$$\int_0^\infty z^{\alpha-1} \mathcal{K}_n(z) dz = 2^{\alpha-2} \Gamma\left(\frac{\alpha - n}{2}\right) \Gamma\left(\frac{\alpha + n}{2}\right)$$

P. Di Bari and A. Riotto : P LB 671, 462 (2009)

Importance of the new parametrization on N2 leptogenesis

Old parametrization:

$$\Omega = \zeta \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos z_{23} & \sin z_{23} \\ 0 & -\sin z_{23} & \cos z_{23} \end{pmatrix} \begin{pmatrix} \cos z_{13} & 0 & \sin z_{13} \\ 0 & 1 & 0 \\ -\sin z_{13} & 0 & \cos z_{13} \end{pmatrix} \begin{pmatrix} \cos z_{12} & \sin z_{12} & 0 \\ -\sin z_{12} & \cos z_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

New parametrization

$$\Omega(z_{12}, z_{13}, z_{23}) = R(\alpha_{12}, \alpha_{13}, \alpha_{23}) \cdot \Omega_{\rm boost}(\vec{\beta}) \,,$$

Asymmetry from N_2 will survive if $K_{i\alpha} < 1$

Randomly generate all the parameters

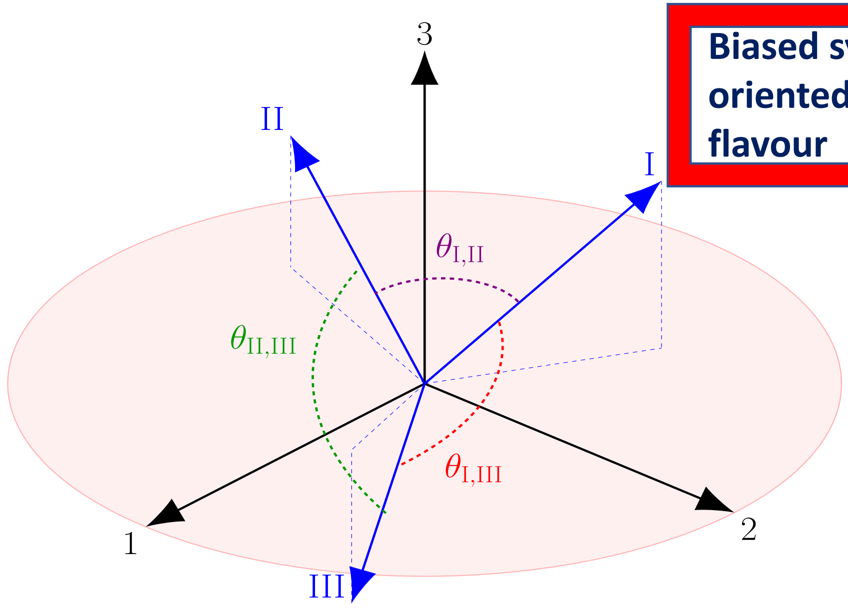
Generate random matrices in a group theoretic way

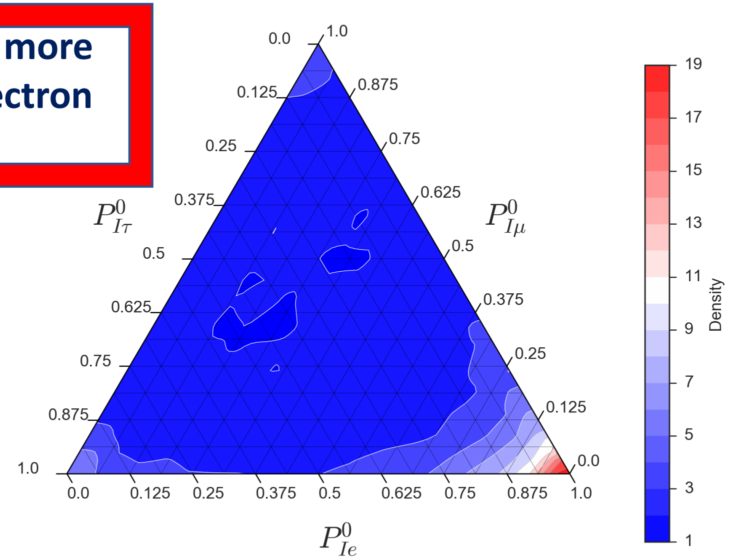
$$K_{I\alpha} = \frac{1}{m^*} \left| \sum_i U_{\alpha i} \sqrt{m_i} \Omega_{iI} \right|^2 ,$$

$$m_{D\alpha J} = U_{\alpha i} \sqrt{m_i}\, \Omega_{iJ} \sqrt{M_J} \,.$$

Generating the decay parameters randomly with no experimental information: all the angles and phases are generated randomly [0,360^0].

Biased system: I is more oriented to the electron flavour





Generating the decay parameters randomly with no experimental information: Using Haar Measure: `Representing seesaw neutrino models and their motions in lepton flavor space', **Rome Samanta**, Pasquale Di Bari and Michele Re Fiorentin. Arxiv: 1812.07720

The leptonic mixing matrix is an element of U(3). Haar Measure corresponding to U(3)

$$dV \equiv d(\sin^2\theta_{12})d(\sin^2\theta_{23})d(\cos^4\theta_{13})\prod_j d\alpha_j,$$

$$K_{I\alpha} = \frac{1}{m^*}\left|\sum_i U_{\alpha i}\sqrt{m_i}\Omega_{iI}\right|^2$$

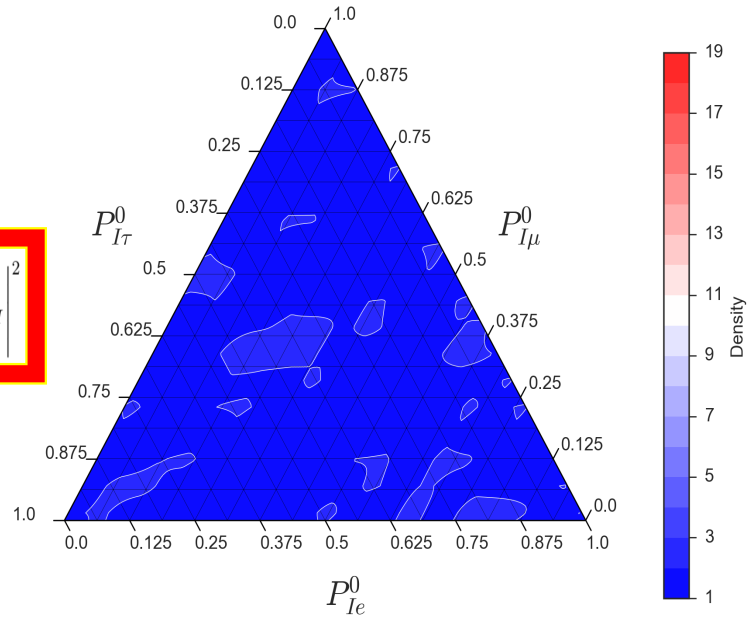
The orthogonal matrix is an element of $SO(3)_C$ which is isomorphic to the Lorentz group $O(3,1)^{+}$.

$$\Omega = \Omega^{\text{Rotation}}\Omega^{\text{Boost}},$$

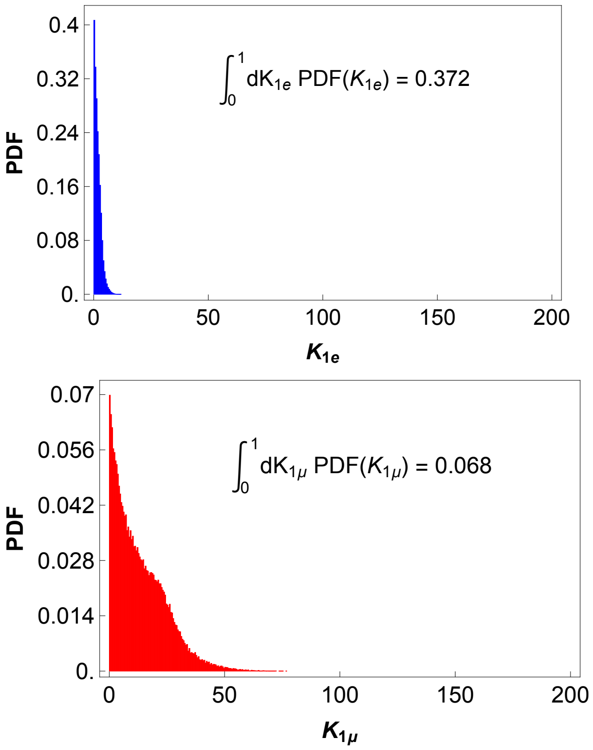
$$\Omega^{\text{Rotation}} \in SO(3)_{\mathbb{R}}$$

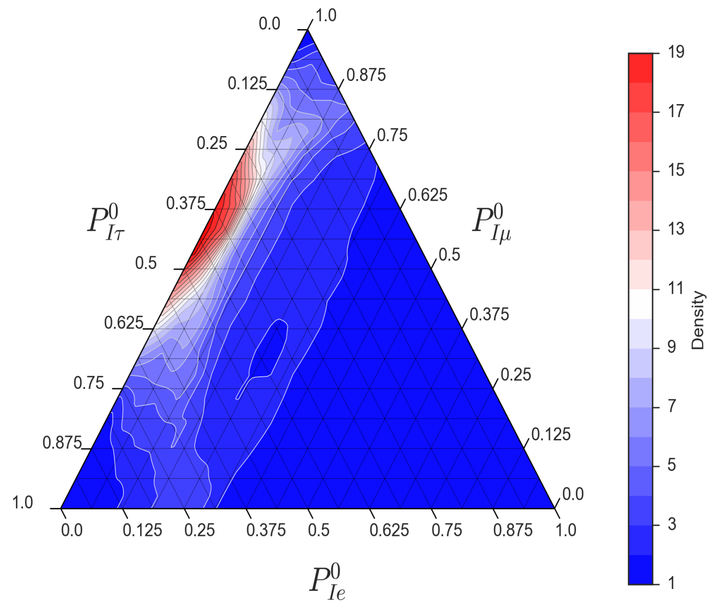
HM

$$dV \equiv d(\sin\phi_2)d\phi_1 d\phi_3.$$

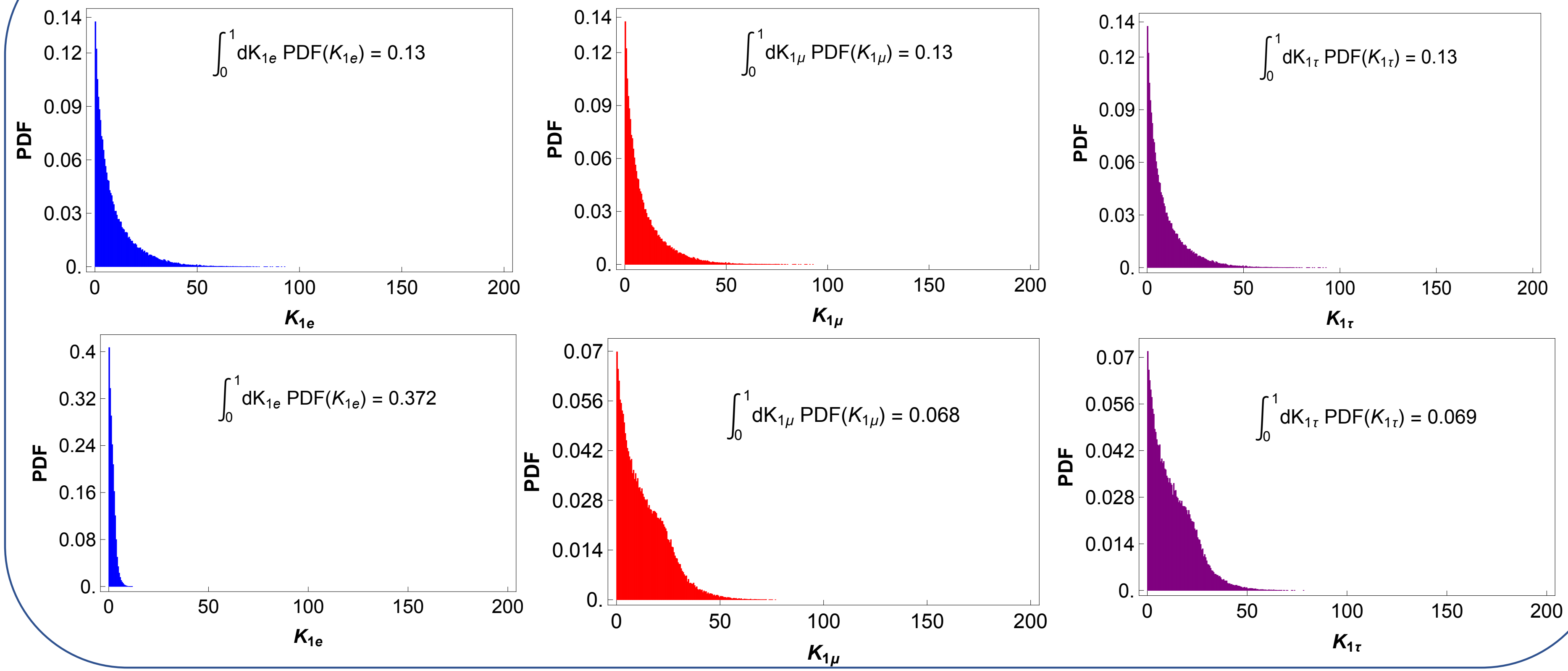


Putting experimental information: NuFiT lateset, 2018

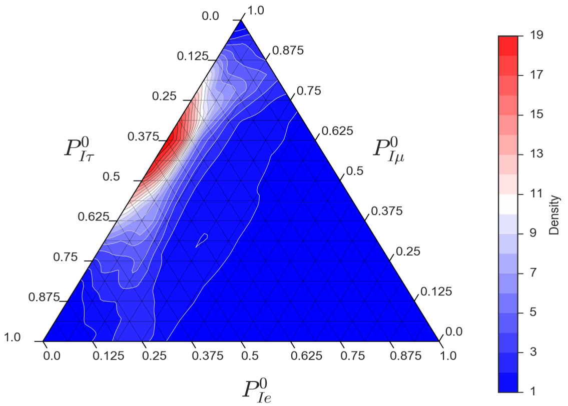


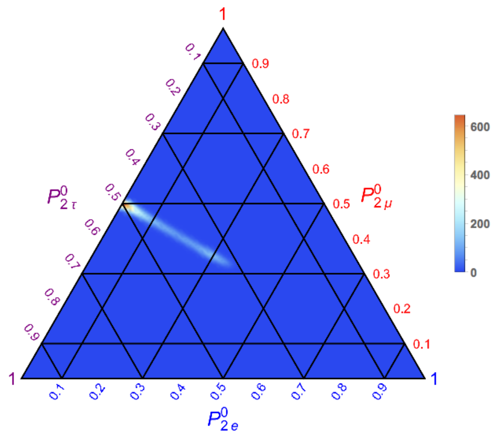


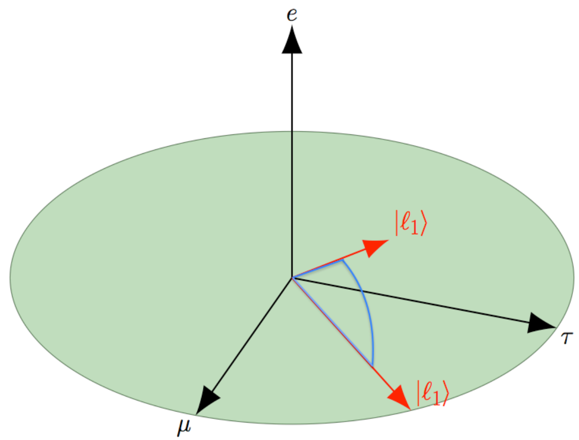
Neutrino oscillation data enhances the probability of the decay parameter being smaller



A model with $CP^{\mu\tau}$ symmetry: $P_{1\mu}= P_{1\tau}$ $\theta_{23}= 45^0$ Cos $\delta = 90^0$ or 270^0







General Case

$CP^{\mu\tau}$ symmetric case

P. Di Bari and Michele Re Fiorentin and RS Arxiv: 1812.07720

E.g., Walter Grimus et al Phys.Lett. B579 (2004)
RS et al JHEP 1806 (2018) 085

Conclusion

1. We have shown how neutrino seesaw model could be visualized graphically

2. We introduce a new matrix called the Bridging matrix (B) that connects the light neutrinos to the heavy neutrinos.

3. We introduce the idea of Lorentz boost in flavour space and show, how this is related to fine tuning in seesaw models.

4. We introduce a new parametrization of the orthogonal matrix and show how this lead to flavour unbiased theory.

5. Neutrino oscillation data creates 'electronic hole' with a higher probability (37%), thus the asymmetry generated by N_2 would more likely to pass through.